

EVOLUTION OF A SIMPLICIAL UNIVERSE

T. Yukawa and S. Horata

The Graduate University of Advanced Studies-SOKENDAI
Hayama, Kanagawa, Japan 240-0193

A. It is the time to bring the quantum gravity view in cosmology.

* Observations of the CMB anisotropy show existence of

i) ***Inflations:***

The anisotropy requires the quantum correlation before the inflation,

and the persistence during the inflation

ii) ***Dark Energy:***

Inflation is the rapid expansion due to the vacuum pressure.

* They should be the consequences of quantized space-time.

B. Quantum Gravity is expected to be to produce the universe.

- *Such a theory should not contain physical laws, since they appear after the creation of space-time.
- * It contains the least assumptions.

An example: *the Peano axioms* for the natural numbers:

1. Existence of the **element** 1 .
2. Existence of the **successor** $S(a)$ of a natural number a .

C. Axiom for the simplicial universe.

1. Existence of the element, a **d-simplex**.
2. Existence of the neighbor to form a **simplicial complex**.

Remarks:

Of course, these conditions are not enough. During the course of creating a universe we might choose additional conditions. In such cases we choose most general or least prejudiced one.

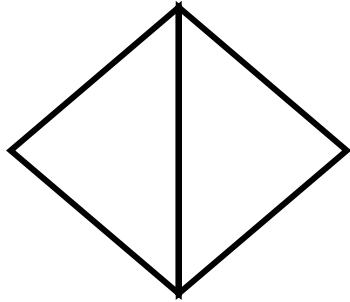
We may think the world has no a priori rules. The rules emerge as a result of the evolution, which sustains large structure with long life.

D. Quick tour of the simplicial quantum gravity.

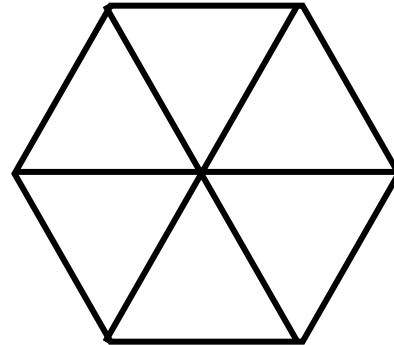
* Let me illustrate the 2d-simplicial quantum gravity as an example.

1. The element = [an equilateral triangle](#)
2. The neighbor = a 2d triangulated surface constructed under [the manifold conditions](#):

i) Two triangles attach through one link.



ii) Triangles sharing one vertex form a disk.



* The partition function is a set of simplicial manifolds with distinct triangulations (configurations).

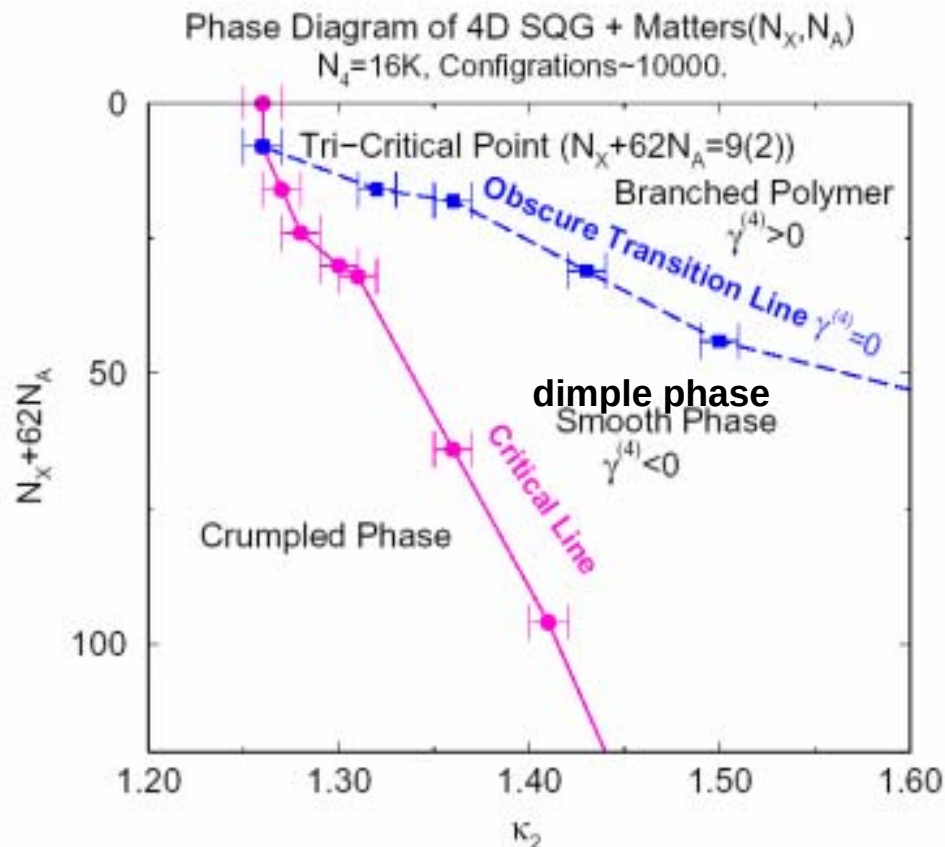
* Known facts:

It is equivalent to [the matrix model](#),
and in the same universality class as [the quantum Liouville theory](#)

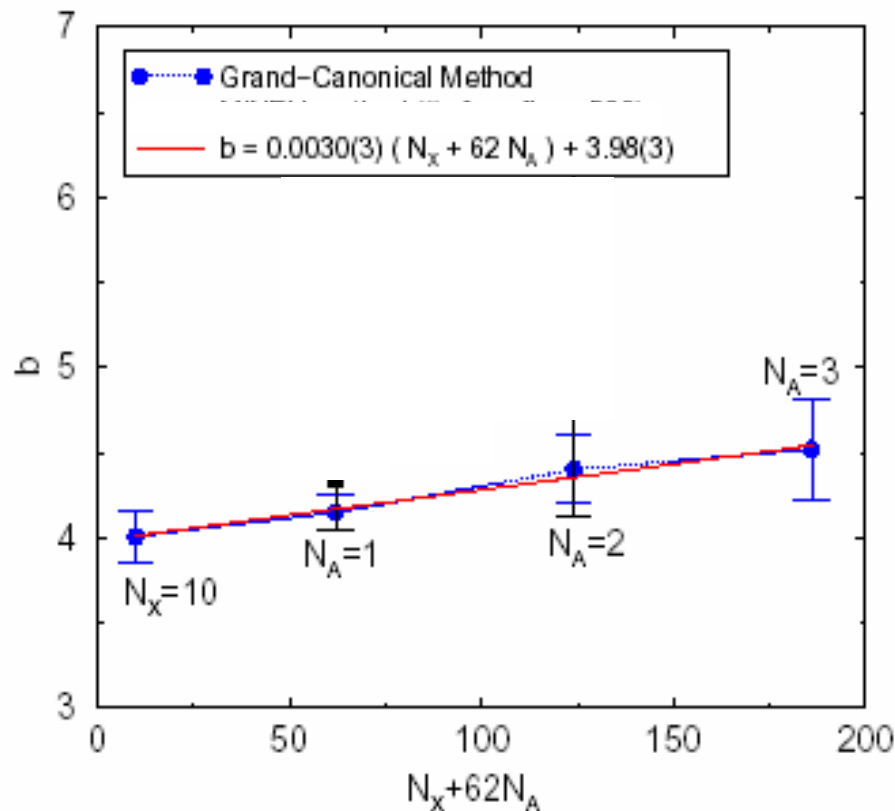
E. What do we know about the 4d-simplicial quantum gravity?

* We have shown numerically:

- i) there exist three phases: *crumple-BP-dimple* , and
- ii) between crumple and dimple phases the transition is continuous as the 2nd order phase transition.



iii) It has the **string susceptibility** similar to the conformal gravity, suggesting to be in the same universality class.



F. Introducing an extra dimension in space with an open boundary.

- * In the standard dynamical triangulation method we used to construct a space with a closed topology, and study its thermo-dynamical property.
- * To describe the evolution we need to consider an ensemble of configurations with varying boundaries, and bring the time coordinate in them.
- * Space with $R \times S^{d-1}$ topology is considered.

G. Let's construct a d -dimensional space with $R \times S^{d-1}$ topology.

* Again, we consider the 2d-dynamical triangulation as an example.

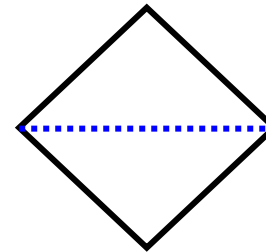
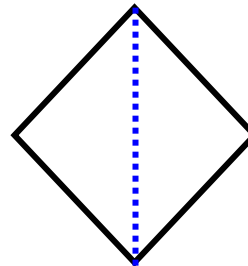
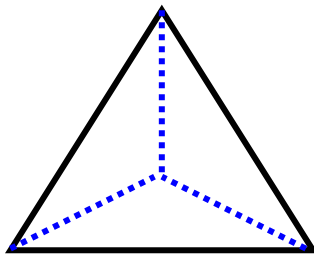
i) *The partition function* consists of a set of spaces with distinct triangulations (configurations) of a fixed S^2 topology weighted appropriately.

ii) Configurations are produced by the *(p,g) -moves* forming a Markov chain:

(1,3)

(3,1)

(2,2)



with the move probability constrained by the detailed balance:

$$\frac{p_a}{n_a} w_{a \rightarrow b} = \frac{p_b}{n_b} w_{b \rightarrow a}$$

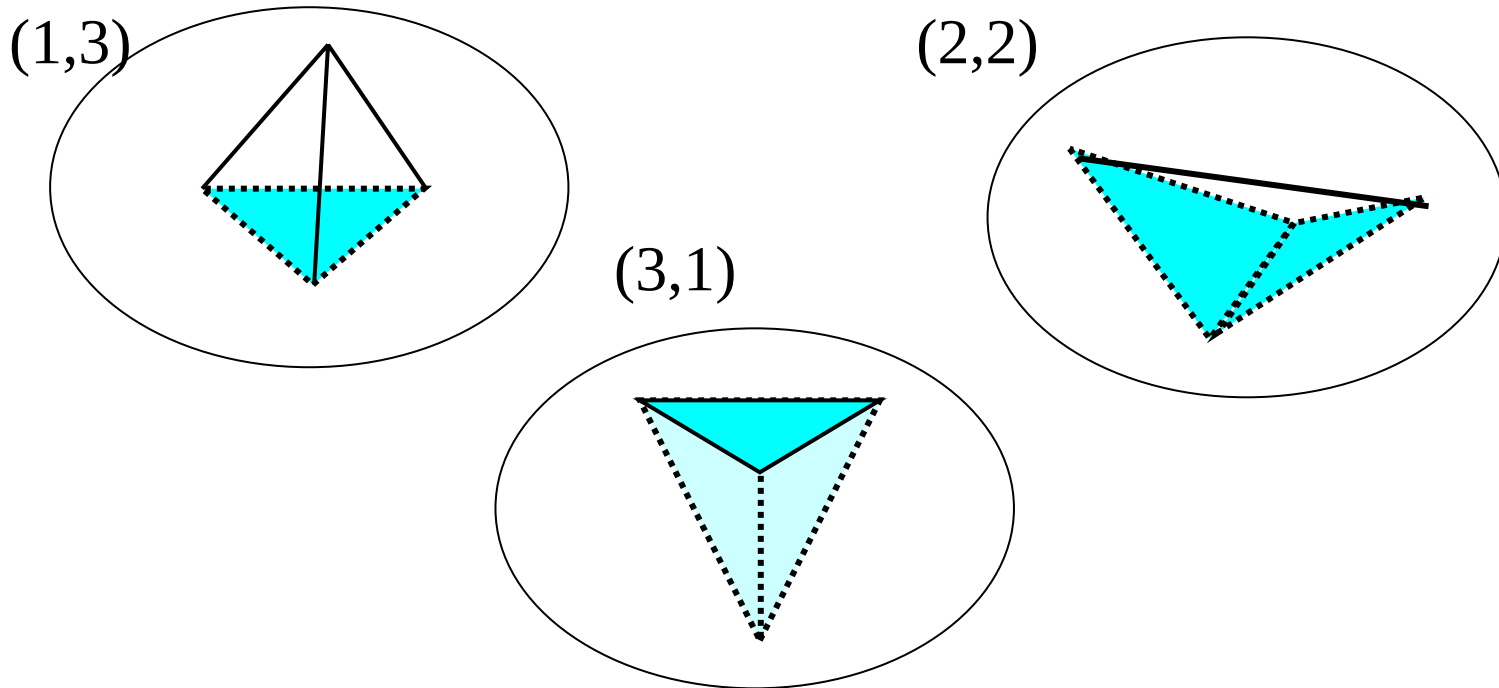
p_a : probability weight for the configuration a ,

n_a : number of all possible moves starting from a configuration a

H. Adding one extra dimension.

*** We regard the S^2 surface as the boundary of a 3-ball(B^3)**

* Moves on the S^2 boundary are regarded by either attaching a tetrahedron on the boundary sharing one, two, or three triangle(s) in common, or to the contrary taking out one tetrahedron:

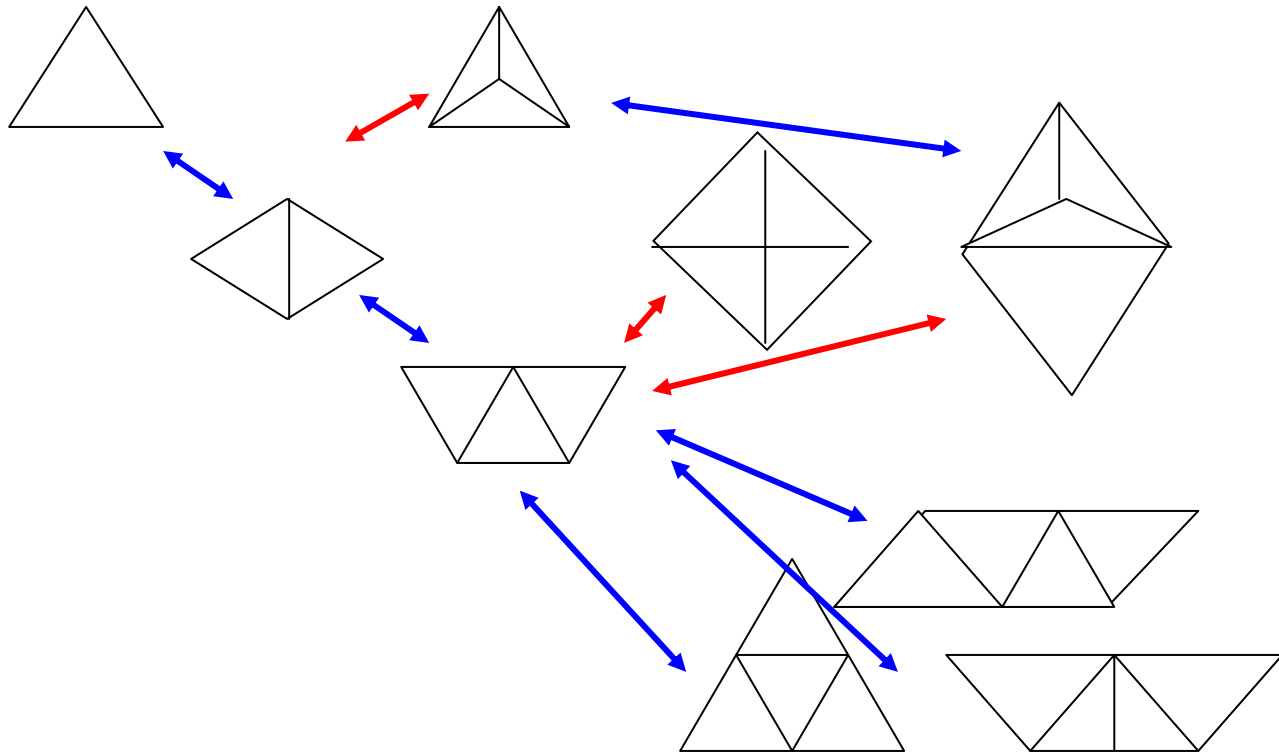


* We call them as $\{\Delta S, \Delta V\}$ -move.

In the 3-d case there are 6 moves: $\Delta S = \{-2, 0, 2\}$ and $\Delta V = \{-1, 1\}$

I. Creation of the quantum universe.

- * The 2d-quantum universe is a set of all possible disks of distinct triangulations with appropriate probabilities.
- * For example up to $N_2=4$:



- * Four types of moves:
 $\{1,1\}, \{-1,1\}$ and the inverses $\{-1,-1\}, \{1,-1\}$.

J. Types of the 2d-universe

- * Number of distinct configurations with the volume V and the area S increases exponentially .
- * There are d control parameters $\{u_i\}$ in d -dimension:

$$A = \sum_{i=0}^d u_i N_i$$

- *Using the Euler relation, $N_2 - N_1 + N_0 = \chi$, which reduces independent parameter by one, for a 2-dimensional where $\{N_i\}$ is the number of i -simplex. universe, we write

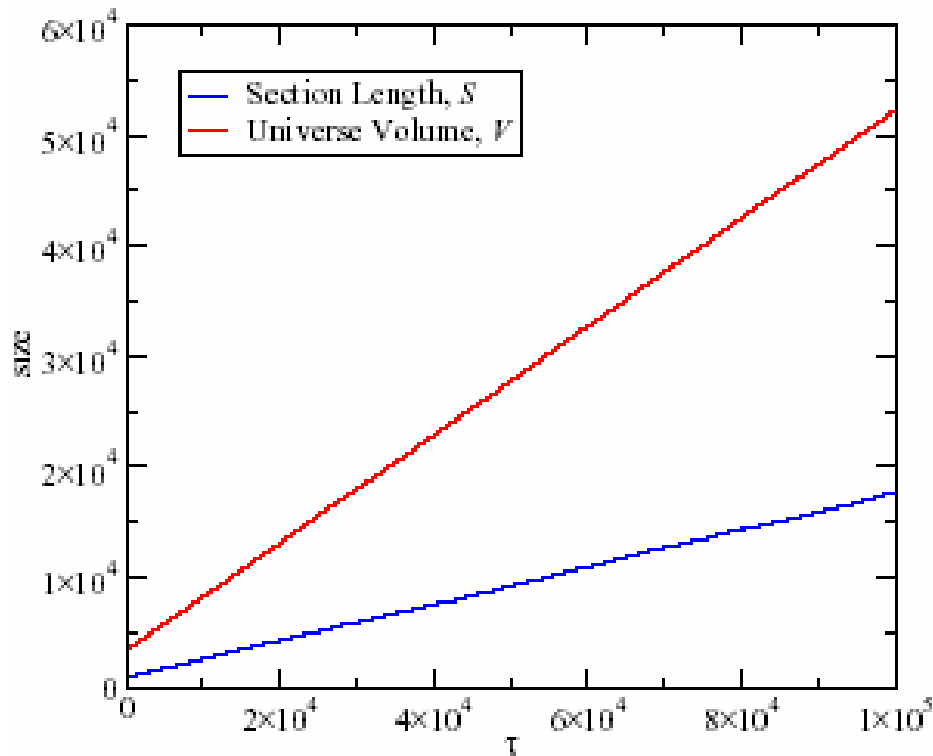
$$A = \mu N_2 + \mu_B \tilde{N}_1$$

where $\tilde{N}_1$ is the number of triangles on the boundary.

- * μ and μ_B are the cosmological constant and the boundary cosmological constant.

Introducing the Monte Carlo computer time τ

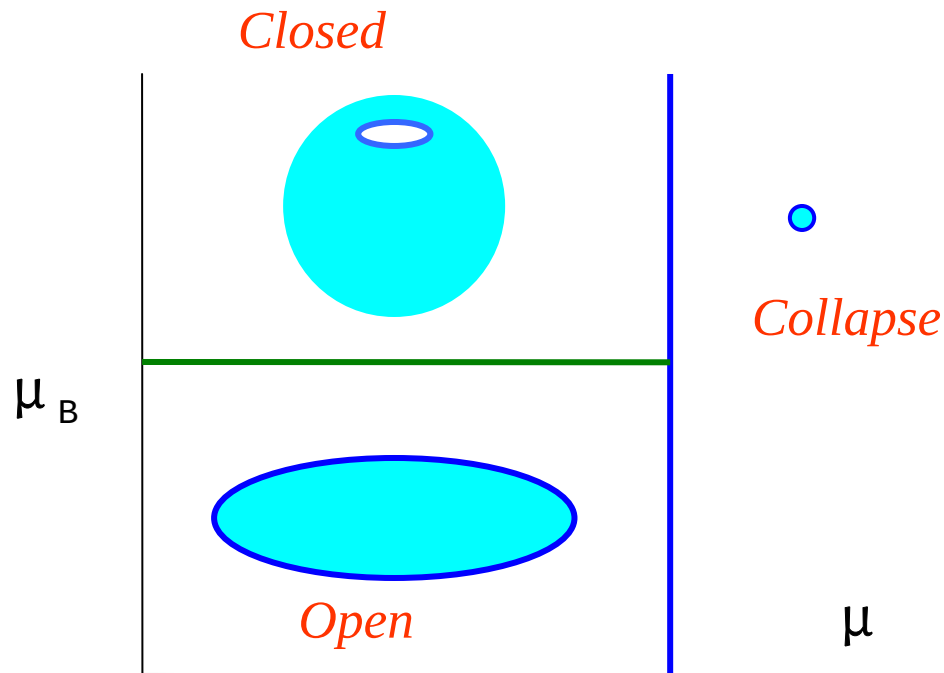
- *Numbers of Metropolis checks in the Monte Carlo method for a Markov chain.
- * The volume V and the area S grow almost linear in τ :



$$V = V_0(\mu, \mu_B)\tau,$$

$$S = S_0(\mu, \mu_B)\tau$$

- * V_0 and S_0 change sign at the critical line μ^c and μ_B^c
- * In the two dimensional parameter space we expect three types of universes: *open-closed-collapse*



L. *Physical time is defined*

* Within a scale factor *the physical time* t is defined through

$$V(t) = \int^t S(t') dt'$$

* These two times are related as

$$t = \int^{\tau} d\tau' \frac{1}{S(\tau')} \frac{dV(\tau')}{d\tau'}$$

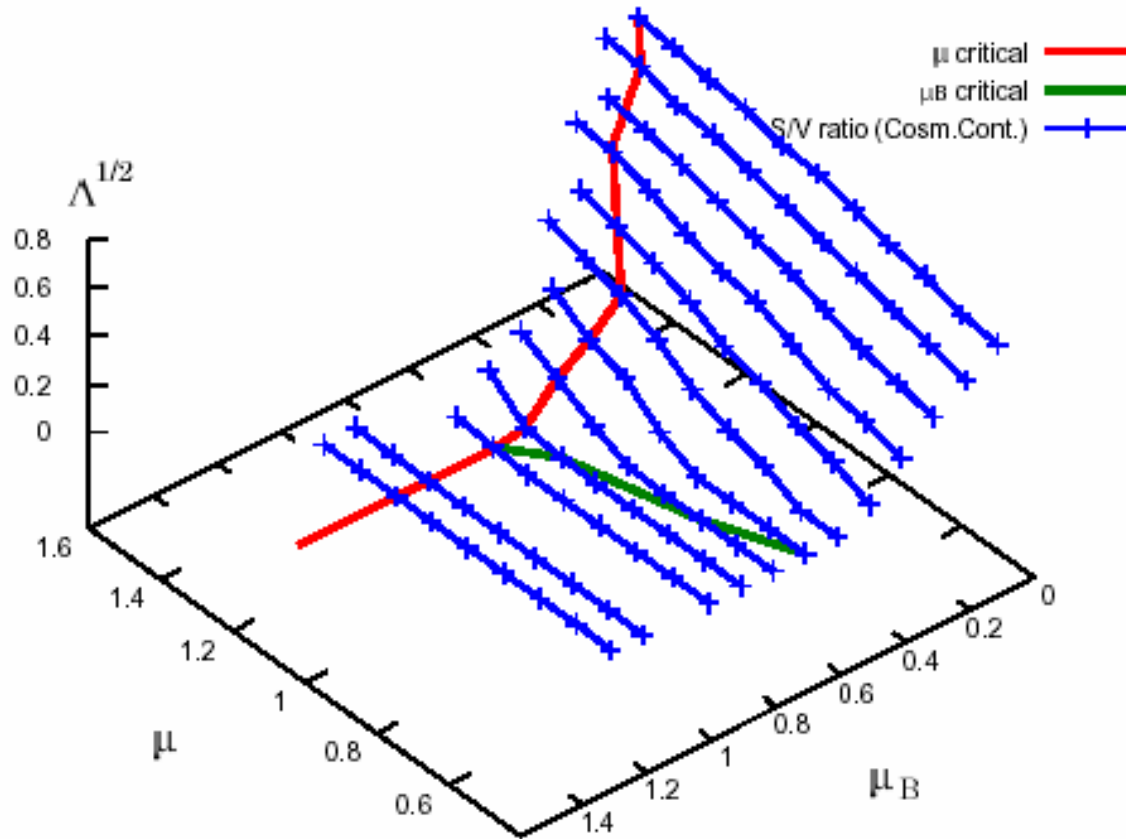
* From the computer simulation

$$\tau \approx \exp\left(\frac{S_0}{V_0} t\right)$$

which means the *inflation!* with the velocity S_0/V_0 .

M. Numerical results

* The inflation velocity $\Lambda^{1/2} = S_0/V_0$:



* The critical values are close to what we expect from the M^3 -matrix model:
 $\mu^c \sim 1.1$ ($1/2 \log[256/27]$)
 at $\mu_B = 0$, (all diagrams)
 $\mu^c \sim 1.5$ ($1/4 \log[432]$)
 at $\mu_B = 0.8$ ($1/2 \log[16/3]$)
 (1PI without tadpole and self-energy diagrams)

N.B. Beyond μ -critical (the red line) the $\Lambda^{1/2}$ changes its sign,
i.e. collapse to a point.

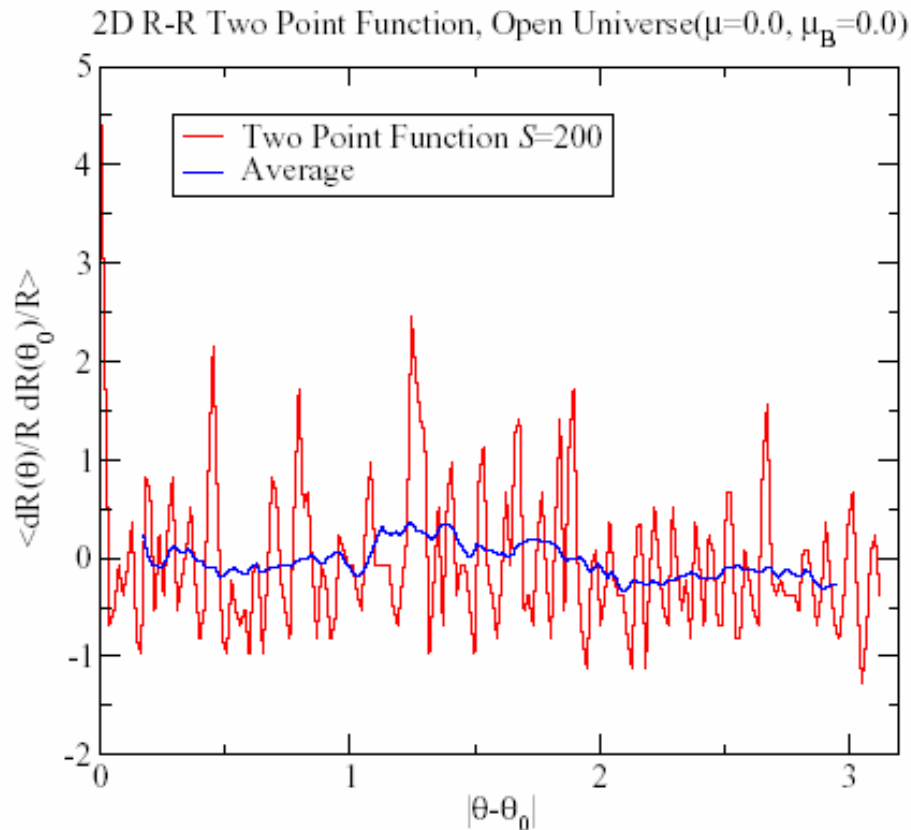
N. Correlations

* Geodesic distance $\sim r \theta$ (r : radius of the universe: $r=S/2 \pi$)

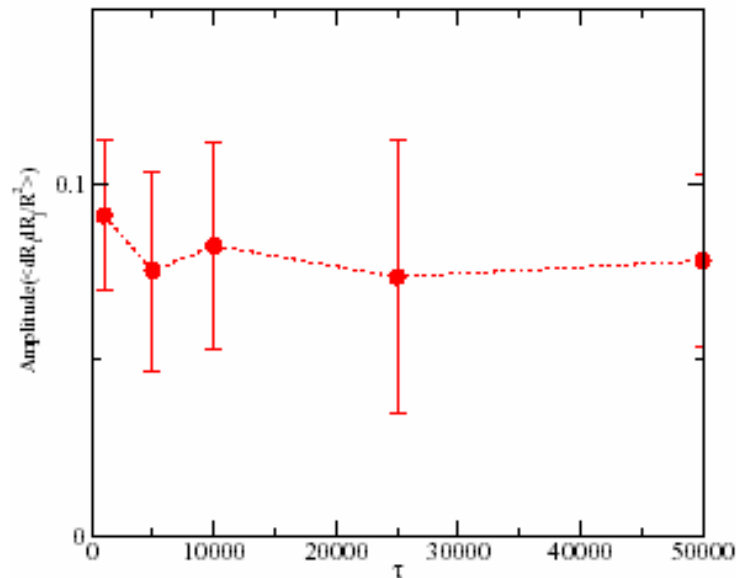
* The two-point correlation function is defined as

$$\left\langle \frac{(R(0) - \langle R \rangle)(R(\theta) - \langle R \rangle)}{\langle R \rangle^2} \right\rangle$$

* Strong quantum correlation at small θ ($\sim 1/r$)

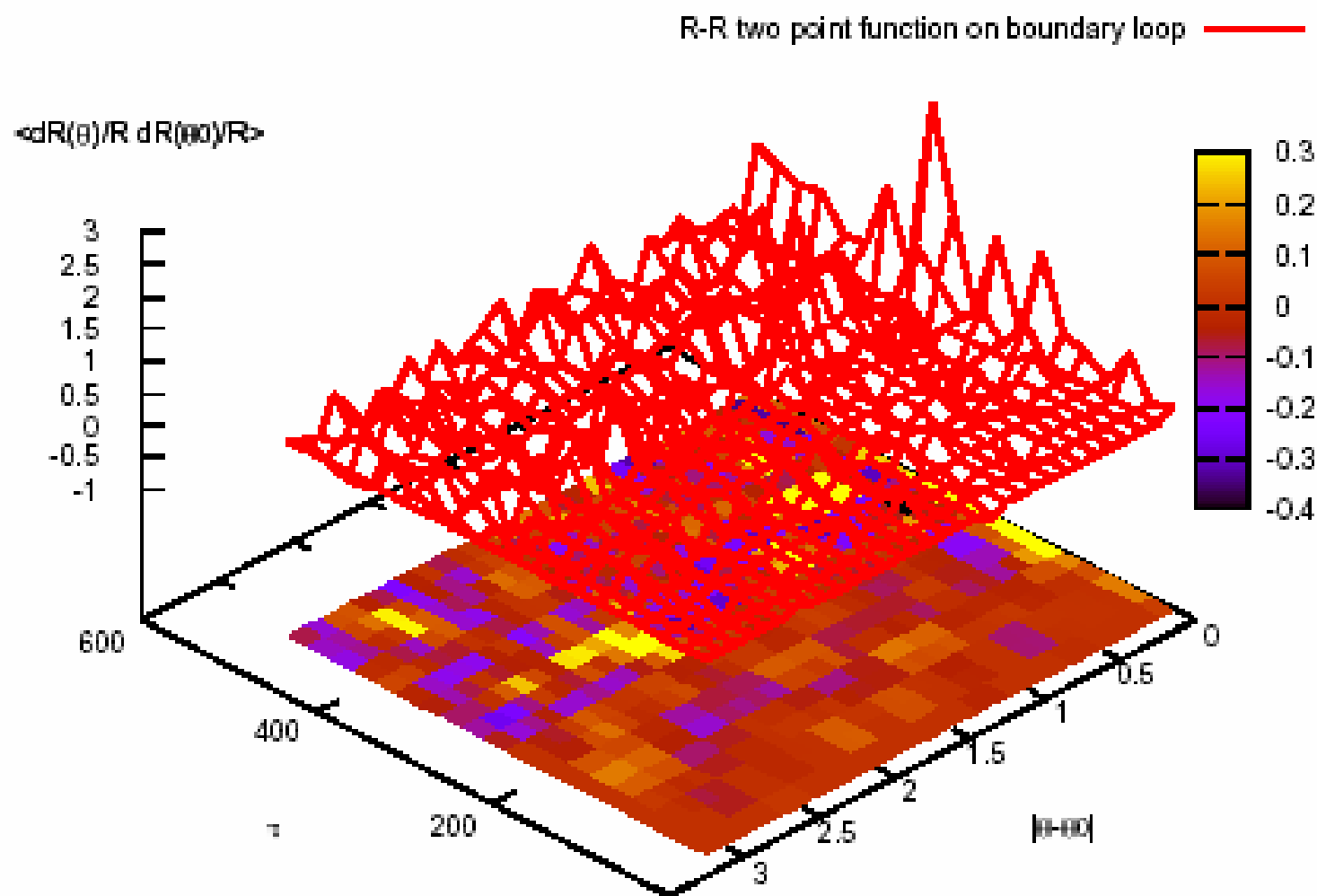


* Large angle correlation remains!

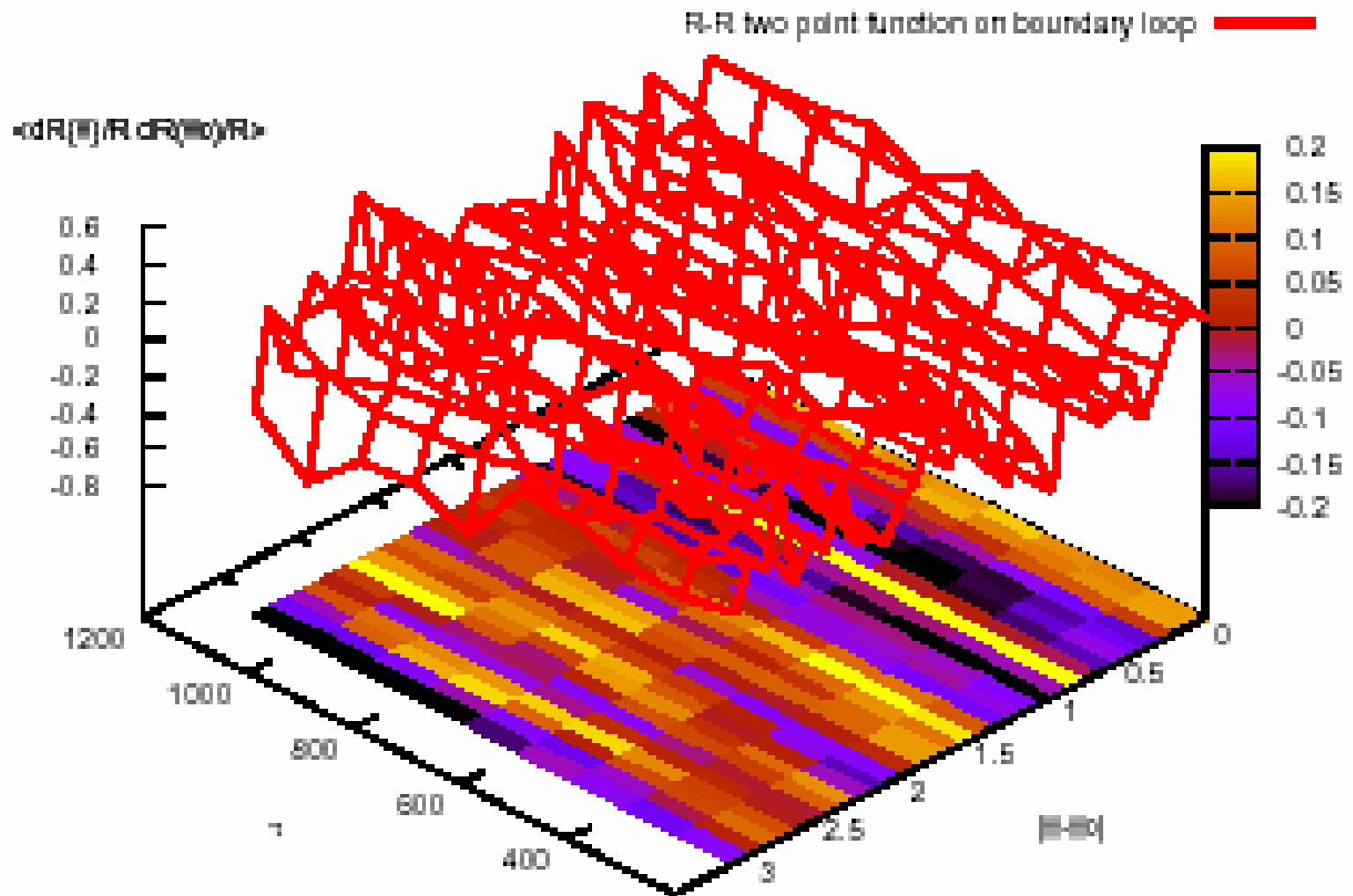


O. Where the large angle correlation is born?

* Initial correlation ($N_s < 500$)



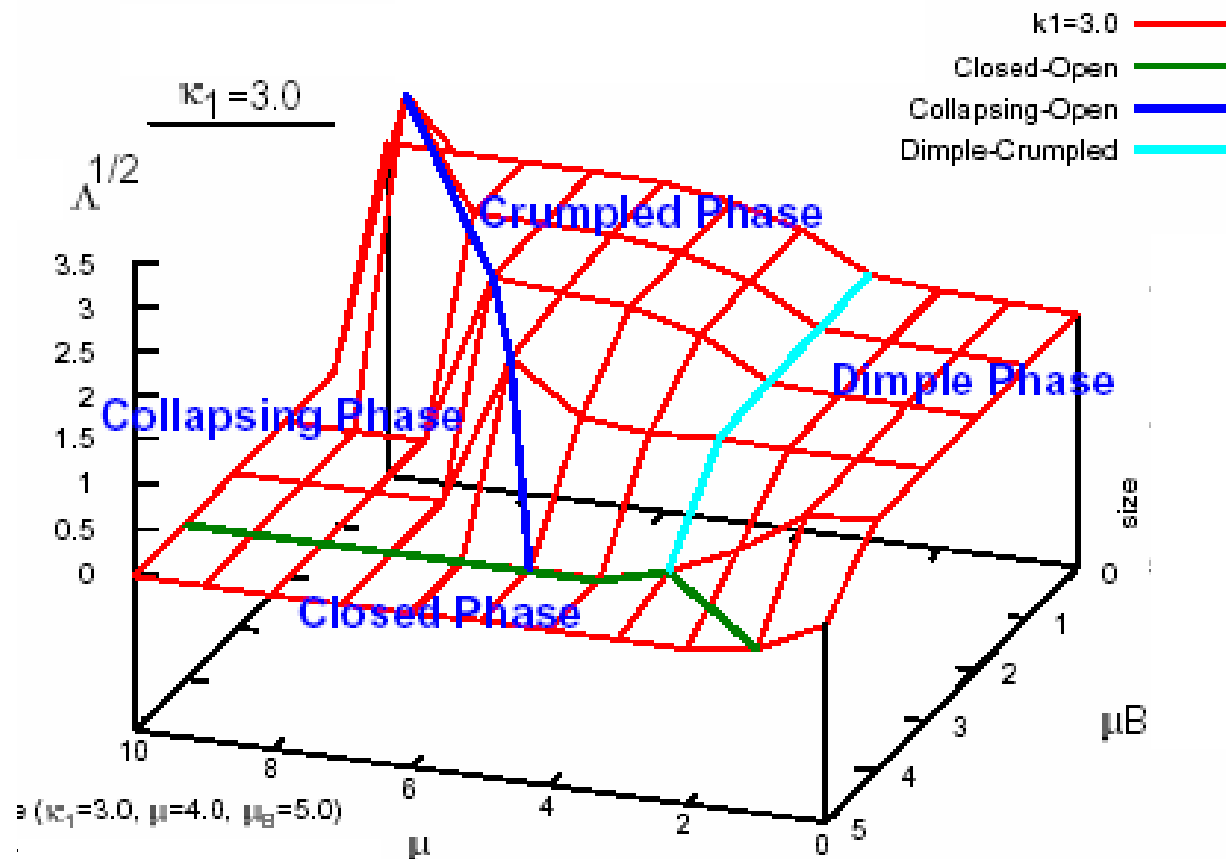
* During the inflation ($N_2 < 1100$)



P. What will come out in higher dimensions?

* In 3d , there are 3 controlling parameters:

$$A = \mu N_3 + \mu_B \tilde{N}_2 - \kappa_1 N_1$$



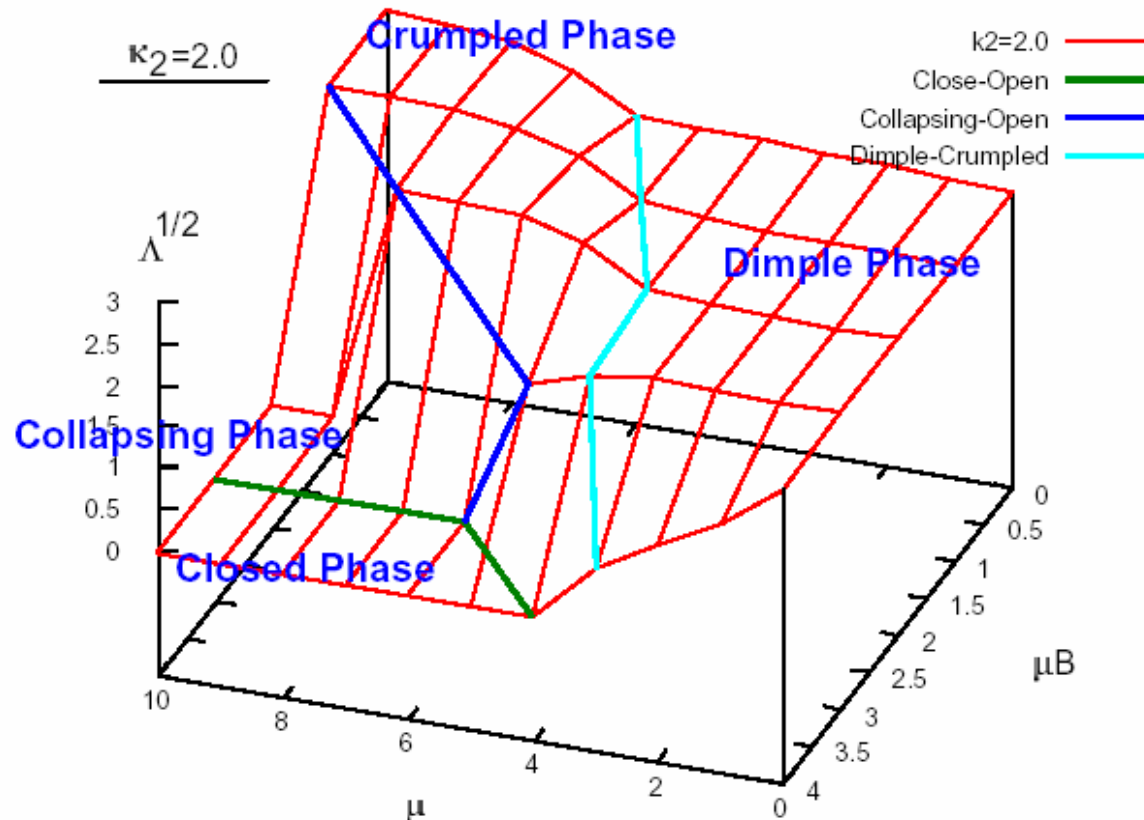
* κ_1 introduces two phases (*dimple-crumple*) in open universe

Q. In 4d

* There are 4 controlling parameters:

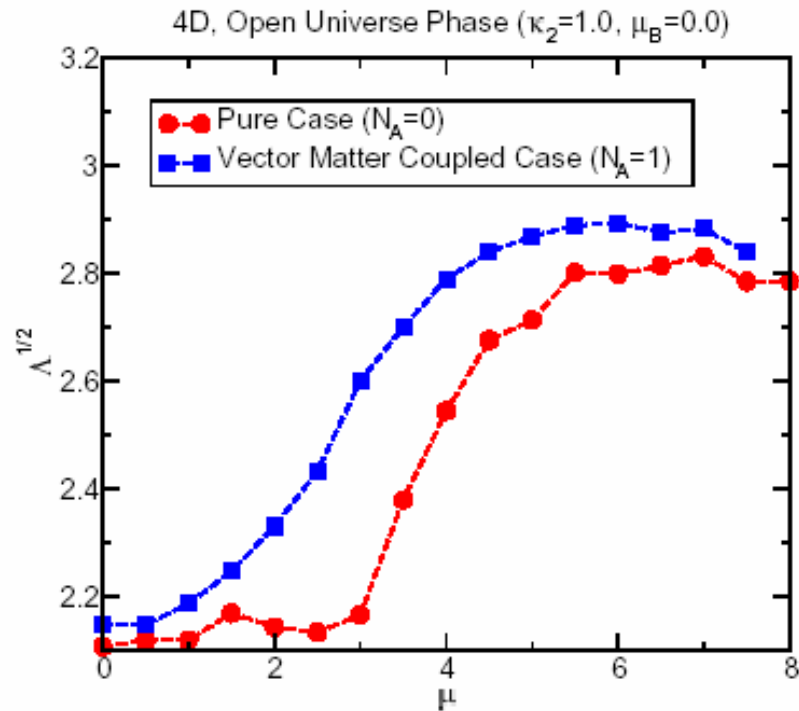
$$A = \mu N_4 + \mu_B \tilde{N}_3 - \kappa_2 N_2 + \kappa_1 N_1$$

* The physical meaning of the new term is not studied, yet.
(Take it to be 0.)



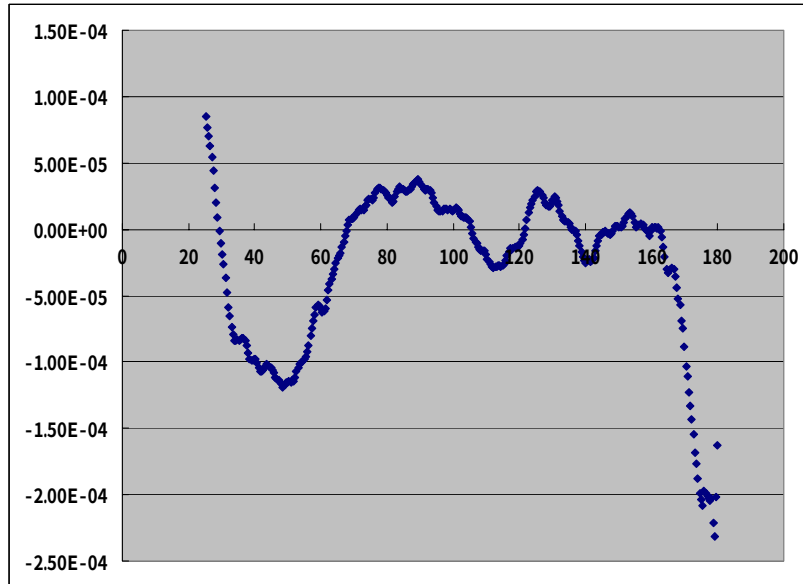
R. Effect of matter fields in 4d

* Adding one vector field the transition between two phases become smooth.



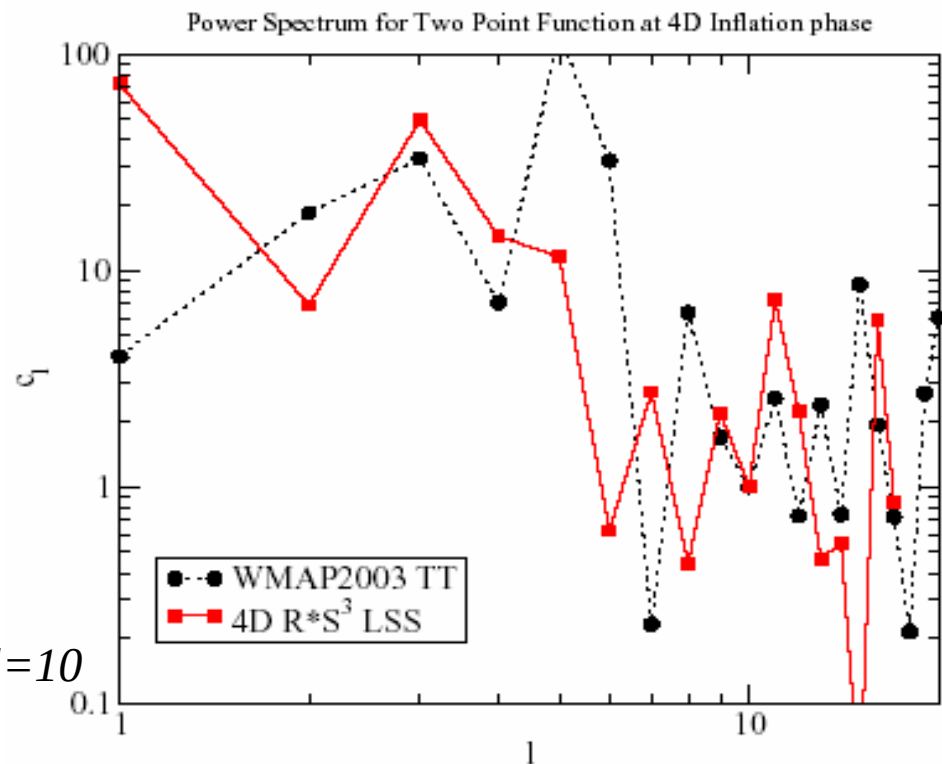
S. How about the correlation in 4d ?(Preliminary)

* Two point correlation on the last scattering surface (lss)
Compared with the WMAP observation (TT-corr)



N.B. Normalized at $l=10$

* The power spectrum in $P_l(\cos \theta)$



To be continued

Thank you for your patience!

